

ICMEs Observed by STEREO A/B

Comparison pair #	#	STEREO	Start time [Year Doy Month/Day HH:MM]	Magnetic obstacle (~ flux rope) start time	End time	Ptmax [pPa]	Bmax [nT]	Vmax [km/s]	ΔV^1 [km/s]	Group ²	Comment
1	1	A & B	2006 348 12/14 14:12	348 12/14 23:04	349 12/15 14:08		18 (20)*				plasma data data gap, like Group 3 event
	2	B	2007 142 5/22 04:20	142 5/22 04:20	142 5/22 22:00	135	17.5	480	-60	1	STEREO B: higher Ptmax, shorter. At both A & B, followed by fast stream
	3	A	2007 142 5/22 14:00	142 5/22 14:00	143 5/23 13:30	63	11.5	540	-80	2	
	4	B	2007 296 10/23 10:35	296 10/23 16:50	297 10/24 00:07	50 (125)	10.3 (10.6)	395 (420)	-35	2	not nice B rotations, slow, low β , followed by a SIR
	5	B	2007 364 12/30 00:30	364 12/30 07:00	at least 365 12/31 23:59	100	12	365 (380)	-85	1	fuzzy plasma data, slow, nice B rotations, left-handed flux rope

()*: values are from the region including the sheath region

ΔV^1 : temporal variation of solar wind speed over one event

Group²: We sort ICMEs into 3 groups depending on their temporal profiles of Pt. Corresponding to the Group 1, 2, and 3 ICMEs, the Pt profile, excluding any shock and/or sheath region (if present), respectively, has a central pressure maximum, a steady plateau, or a gradual decay. In the hypothesis that all ICMEs have a central flux rope, these three groups of Pt profiles are due to different approach distances to the central flux rope. Group 1 ICMEs are assumed to be the ones penetrated by spacecraft near the flux rope axis, and they usually present signatures of magnetic clouds. See Jian et al. [2006a] for more detail.